How Radial Oscillations Can Help to Probe the Onset of the Deconfinement Phase Transition and Special Points in Hybrid Stars

<u>Christoph Gärtlein\*,</u> David Blaschke, Oleksii Ivanytskyi, Violetta Sagun





University of Lisbon (IST) University of Coimbra University of Wroclaw

Fundação para a Ciênci e a Tecnologi 63. Cracow School of Theoretical Physics

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Outline

Compact Stars Building up an NS Special points Radial Oscillations Conclusions

# Outline

Compact Stars

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Conclusions

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# **Compact Stars**





## **Neutron Stars**

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	Neutron star	White dwarf	Sun
$M_{max}(M_{\odot})$	2	1.44	1
R (km)	11-12	10 <sup>4</sup>	$7 \cdot 10^{5}$
$n_c (g/cm^3)$	$10^{14} - 10^{15}$	$10^{7}$	$10^{2}$
rotation speed $(s)$	$10^{-3} - 1$	100	$2 \cdot 10^{6}$
B (G)	$10^8 - 10^{16}$	100	1
T (K)	$10^6 - 10^{11}$	$10^{3}$	10 <sup>5</sup>

Δ

## Variety in the composition

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**Compact Stars** 

# **Construction the Equation of State (EoS)**

#### **Idea:**

combine the EoS for <u>hadronic matter</u> and our model for <u>quark matter</u> via Maxwell construction

Shahrbaf, Blaschke+ (2022)

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DD2npY-T model (hadronic) neutrons and

protons + hyperonic degrees of freedom **Quark matter:** Ivanytskyi,Blaschke (2022) confining relativistic density functional approach (RDF) Underlying Lagrangian:

 $\mathcal{L} = \mathcal{L}_{\textit{free}} + \mathcal{L}_{\textit{V}} + \mathcal{L}_{\textit{D}} - \mathcal{U}$ 

#### Meaning what?

ην Including <u>vector</u> <u>repulsion</u> and <u>diquark</u> <u>pairing</u> controlled by dimensionless couplings

Confinement + Conformal limit + Color Superconductivity

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## **Hybrid EoS**



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#### Building up an NS



#### Building up an NS

# **Special Points and Onset Mass**

- M-R curves show many interesting characteristics
- Plotting curves characterized by the couplings of the quark phase
- Special points and onset mass (deconfinement phase transition)

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#### Special points



Special points







- Similar to stars (well studied for the Sun) radial and nonradial oscillations are considerable
- Changes of radius and pressure (harmonic oscillations)
- Radial: n=0 --> zero or fundamental mode (f-mode)
- n > 1: p-modes

#### **Pressure deviation**



#### **Radial Oscillations**

## How do they work?

### **Equations:**

$$\frac{d\xi}{dr} = -\left(\frac{3}{r} + \frac{1}{\epsilon + p}\frac{dp}{dr}\right)\xi - \frac{\eta}{r\gamma}, \qquad \xi \equiv \frac{\Delta r}{r} \quad \text{radial perturbations}$$

$$\frac{d\eta}{dr} = \omega^2 \left[\frac{\epsilon + p}{p}re^{(\lambda - \nu)}\right]\xi$$

$$-\left[\frac{4dp}{pdr} + 8\pi(\epsilon + p)re^{\lambda} - \frac{r}{p(\epsilon + p)}\left(\frac{dp}{dr}\right)^2\right]\xi$$

$$-\left[\frac{\epsilon}{p(\epsilon + p)}\frac{dp}{dr} + 4\pi\zeta re^{\lambda}\right]\eta, \qquad \eta \equiv \frac{\Delta p}{p} \quad \text{pressure perturbations}$$

- radial pressure oscillations (p-waves) ⇒ no GW emission
- could couple to nonradial oscillations  $\Rightarrow$  emit GW
- used to probe composition of compact stars



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#### Radial Oscillations



- **1.** Phenomenological EoS in agreement with astrophysical constraints
- 2. Variation of  $\eta_D$  while keeping  $\eta_V$  fixed  $\longrightarrow$  family of hybrid quark-hadron EoS also other possibilities corresponding mass-radius curves intersect in **SP**
- **3.** Finding special points within microscopic approach

4. Study of f-modes (belonging to SPs)

for same gravitational mass, radius and vector coupling

frequency decreases with decreasing diquark coupling

**5.** Earlier deconfinement phase transition **—** higher frequency

6. Observing low f-modes —> closer to maximum mass

with relation: constraining NS mass at which deconfinement transition occurs

CD	$M_{SP}$	R <sub>SP</sub>	$\eta_V$	$\eta_D$	$M_{onset}$	$[M_{max}]$	f
51	$[M_{\odot}]$	[km]			$[M_{\odot}]$	$[M_{\odot}]$	[kHz]
nlq 1.973 1	11.06	0.230	0.749	0.251	2.044	0.840	
			0.740	0.506	2.011	0.821	
			0.731	0.826	1.986	0.770	
			0.721	1.169	1.974	0.318	
			0.711	1.483	1.976	_	
magenta 5.095	11.46	0.290	0.760	0.251	2.159	0.760	
			0.753	0.506	2.130	0.755	
			0.745	0.826	2.104	0.750	
			0.737	1.169	2.094	0.410	
				0.730	1.483	2.095	_
				A 1990	O OF 4	0 0 0 F	a = aa

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#### **Outlook and further findings**

- Constraining the possible values of couplings
- range suggests an early deconfinement of quark matter and deconfinement onset masses below 1 solar mass
- consistent with the existence of heavy NSs with up to 2.4 solar masses



# Thank you for your attention

Any Questions?



End Of Presentation

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Lagrangian of the quark matter theory: including Up and Down quarks

$$\mathcal{L} = \overline{q}(i\partial \!\!\!/ - m)q + \mathcal{L}_V + \mathcal{L}_D - \mathcal{U},$$

$$\mathcal{L}_{V} = -G_{V}(\overline{q}\gamma_{\mu}q)^{2} + \Theta_{V},$$
  
$$\mathcal{L}_{D} = G_{D}(\overline{q}i\gamma_{5}\tau_{2}\lambda_{A}q^{c})(\overline{q}^{c}i\gamma_{5}\tau_{2}\lambda_{A}q) - \Theta_{D}.$$

 $\mathcal{U} = D_0 \left[ (1+\alpha) \langle \overline{q}q \rangle_0^2 - (\overline{q}q)^2 - (\overline{q}i\gamma_5 \vec{\tau}q)^2 \right]^{\frac{1}{3}}$ 

#### **Definition of couplings:**

$$\begin{split} G_V &= \frac{G_{V0}}{1 + \frac{8}{9M_g^2} \left(\frac{\pi^2 \langle q^+ q \rangle}{2}\right)^{2/3}}, \\ G_D &= \frac{G_{D0}}{1 + \frac{8}{9M_g^2} \left(\frac{\pi^2 |\langle \overline{q}^c i \tau_2 \gamma_5 \lambda_2 q \rangle|}{2}\right)^{2/3}}, \\ \eta_V &\equiv \frac{G_{V0}}{G_{S0}} \quad \text{and} \quad \eta_D \equiv \frac{G_{D0}}{G_{S0}} \quad \text{Vacuum value of scalar coupling} \\ \text{dimensionless} \end{split}$$

Fixing parameters via mass and decay constant of the pion as well as of the scalar meson

$m \; [MeV]$	$\Lambda \ [MeV]$	$\alpha$	$D_0 \Lambda^{-2}$	$M_g$ [MeV]
4.2	573	1.43	1.39	600

#### Speed of sound profiles

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